Searching for dark energy off the beaten track

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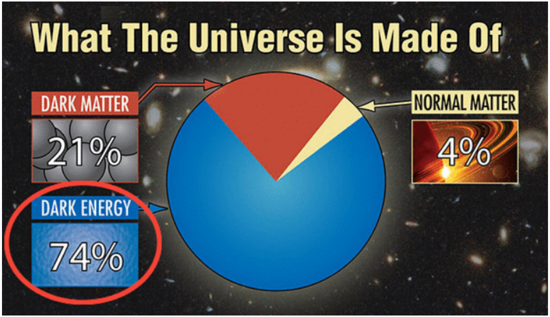
Friday Cosmology Seminar
Szczecińska Grupa Kosmologiczna, 11 March 2022

Dark Energy



- Part I: direct detection of DE on Earth
- Part II: consistency tests of ΛCDM, implications for (early and late) DE
- Part III: new ways to search for light particles (related or not to DE?)

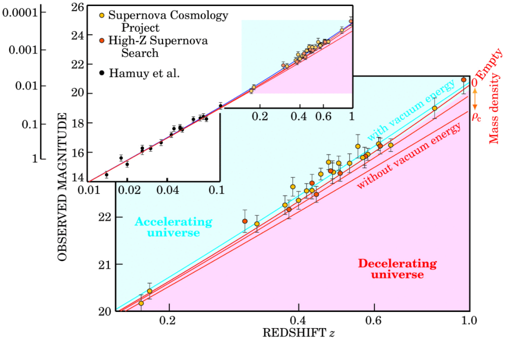
Note: blue $\rightarrow$ (Master's/PhD) students, red $\rightarrow$ postdocs

Student's name (student's institution)

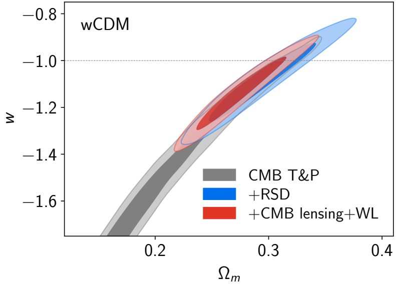
Postdoc's name (postdoc's institution)

The beaten track

Gravitational signatures of DE: the effect of DE's energy density on the background expansion or the growth of structure, probed by standard cosmological observations, with particular focus on DE's equation of state $w_{\rm DE} = P_{\rm DE}/\rho_{\rm DE}$ $(\sim -1?)$



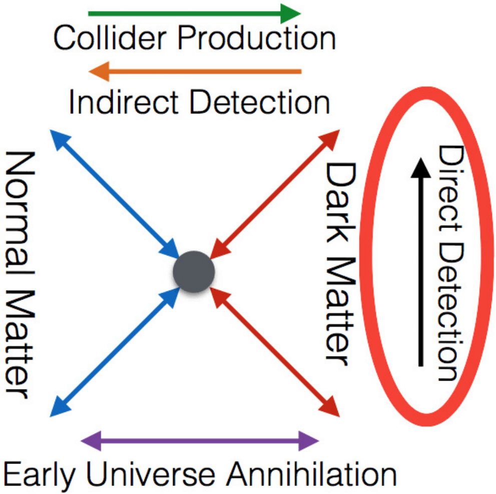
Credits: Perlmutter, Physics Today 56 (2003) 53



eBOSS collaboration, PRD 103 (2021) 083533

Part I: direct detection of dark energy

Are gravitational signatures all there is?



What about dark energy?

Credits: (adapted from) Matt Buckley

Can dark energy and visible matter talk to each other?

Quintessence and the Rest of the World: Suppressing Long-Range Interactions

Sean M. Carroll
Phys. Rev. Lett. **81**, 3067 – Published 12 October 1998

If DE due to a new particle, this typically will:

- be very light [$m \sim H_0 \sim \mathcal{O}(10^{-33})\,\mathrm{eV}$]
- have gravitational-strength coupling to matter

Result/immediate obstacle: long-range fifth forces!

$$F_5 = -\frac{1}{M_5^2}\frac{m_1 m_2}{r^2}e^{-r/\lambda_5}\,, \quad M_5 \sim M_{\rm Pl}\,, \quad \lambda_5 \sim m^{-1} \sim H_0^{-1}$$

Screening

How to satisfy fifth-force tests?

- Tune the coupling to be extremely weak $[M \gg M_{\rm Pl}]$
- Tune the range to be extremely short $[\lambda \ll \mathcal{O}(\rm mm)]$
- Tune the dynamics so the force weakens based on its environment $\longrightarrow$ **screening**!

(At least) 3 ways to screen

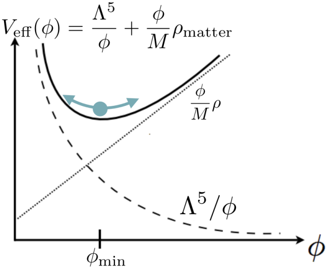
$$F_5 = -\frac{1}{M_5^2(x)} \frac{m_1 m_2}{r^{2-n(x)}} e^{-r/\lambda_5(x)}$$

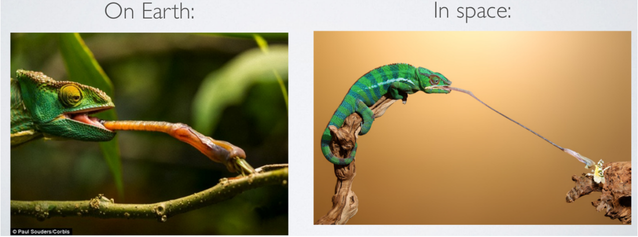
- $\lambda_5(x) \rightarrow$ **chameleon** screening (short range in dense environments)
- $M_5(x) \rightarrow$ symmetron screening (weak coupling in dense environments)
- $n(x) \rightarrow$ Vainshtein (force drops faster than $1/r^2$ around objects)

Chameleon screening

Fifth force range $\lambda(x)$ becomes short in dense environments, scalar field minimizes effective potential determined by coupling to matter

$$
\begin{aligned}
V_{\text{eff}} &= V(\phi) + \phi\rho_m/M \\
m_{\text{eff}}^2 &= \frac{d^2 V_{\text{eff}}}{d\phi^2}\Big|_{\phi=\phi_{\text{min}}} \propto \rho^n \,, n > 0 \\
\lambda &\sim 1/m_{\text{eff}} \propto \rho^{-n/2}
\end{aligned}
$$





Credits: Ben Elder

Direct detection of dark energy

Can we detect (screened) DE in DM direct detection experiments?

NATIONAL GEOGRAPHIC POLSKA — NATIONAL GEOGRAPHIC — TRAVELER

LUDZIE

Ciemna energia istnieje? We Włoszech zaobserwowano tajemnicze zjawisko

PHYSICAL REVIEW D **104**, 063023 (2021)

Direct detection of dark energy: The XENON1T excess and future prospects

Sunny Vagnozzi,[1,2,*,¶] Luca Visinelli,[3,4,5,†,¶] Philippe Brax,[6,‡] Anne-Christine Davis,[7,1,§] and Jeremy Sakstein[8,∥]

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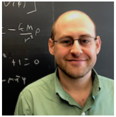
Luca Visinelli (Shanghai)


Phil Brax (IPhT, Saclay)


Anne Davis (Cambridge)

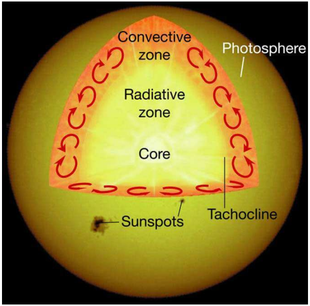

Jeremy Sakstein (Hawaii)

Direct detection of dark energy

Production

$$\mathcal{L}_{\phi\gamma} \supset \underbrace{-\beta_\gamma \frac{\phi}{M_{\rm Pl}} F_{\mu\nu}F^{\mu\nu}}_{\text{(anomalous)}} + \underbrace{\frac{T^{\mu\nu}_\gamma \partial_\mu\phi\partial_\nu\phi}{M_\gamma^4}}_{\text{disformal}}$$

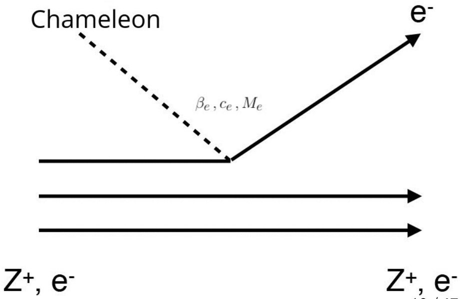
Production in strong magnetic fields of the tachocline



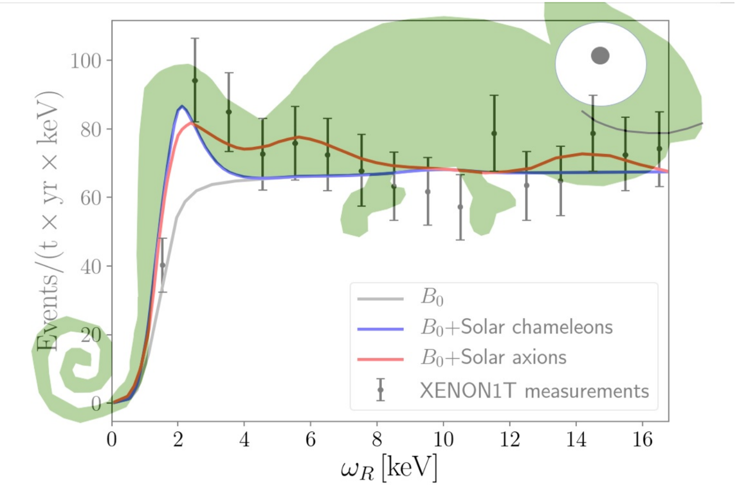
Detection

$$\mathcal{L}_{\phi i} \supset \underbrace{\beta_i \frac{\phi T_i}{M_{\rm Pl}}}_{\text{conformal}} - \underbrace{c_i \frac{\partial^\mu\phi\partial_\mu\phi}{M^4}}_{\text{kinetic-conformal}} T_i + \underbrace{\frac{T^{\mu\nu}_i \partial_\mu\phi\partial_\nu\phi}{M_i^4}}_{\text{disformal}}$$

Analogous to photoelectric and axioelectric effects



Direct detection of (chameleon-screened) dark energy



SV *et al.*, PRD 104 (2021) 063023 Image editing credits: Cristina Ghirardini

Cosmological direct detection of dark energy

Wouldn't scattering between DE and baryons mess up cosmology?

Monthly Notices of the ROYAL ASTRONOMICAL SOCIETY

MNRAS **493**, 1139–1152 (2020) doi:10.1093/mnras/staa311
Advance Access publication 2020 February 3

Do we have any hope of detecting scattering between dark energy and baryons through cosmology?

Sunny Vagnozzi [1★†], Luca Visinelli,[2] Olga Mena[3] and David F. Mota[4]

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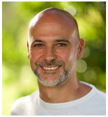
Surprisingly not!

Luca Visinelli (Shanghai)

Olga Mena (Valencia)



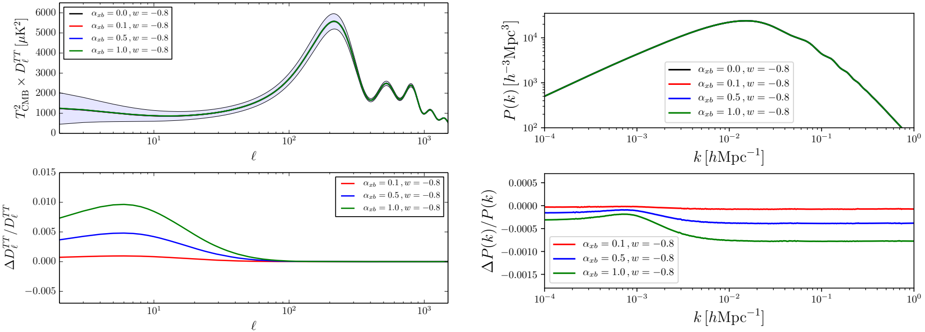
David Mota (Oslo)

Cosmological direct detection of dark energy?

$$\dot{\theta}_b = -\mathcal{H}\theta_b + c_s^2 k^2 \delta_b + \frac{4\rho_\gamma}{3\rho_b} a n_e \sigma_T (\theta_\gamma - \theta_b) + (1 + w_x)\frac{\rho_x}{\rho_b} a n_e \sigma_{xb} (\theta_x - \theta_b)$$

$$\dot{\theta}_x = -\mathcal{H}(1 - 3c_s^2)\theta_x + \frac{c_s^2 k^2}{1 + w_x}\delta_x + a n_e \sigma_{xb}(\theta_b - \theta_x)$$

Impact on CMB and *linear* matter power spectrum ($\alpha = \sigma_{xb}/\sigma_T$)



SV *et al.*, MNRAS 493 (2020) 1139

N-body simulations of DE-baryon scattering

What about the non-linear regime?

Cosmological direct detection of dark energy: non-linear structure formation signatures of dark energy scattering with visible matter

Fulvio Ferlito,[1,2★] Sunny Vagnozzi,[3†‡] David F. Mota[4] and Marco Baldi[2,5,6]
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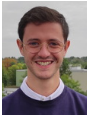
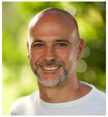
Only one way to find out: run N-body simulations!



Fulvio Ferlito (MPA Garching)

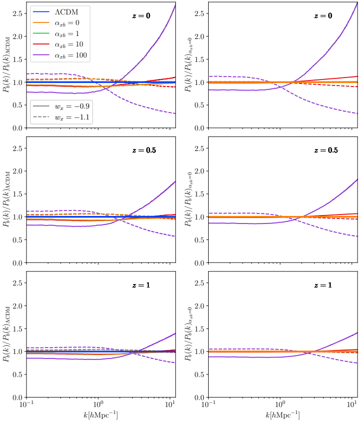
Marco Baldi (Bologna)



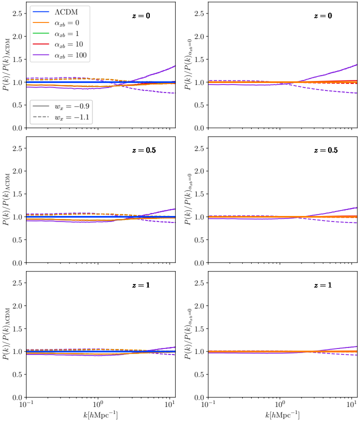
David Mota (Oslo)

N-body simulations of DE-baryon scattering

Baryon power spectrum relative to ΛCDM (left) and no-scattering wCDM (right)



Matter power spectrum relative to ΛCDM (left) and no-scattering wCDM (right)



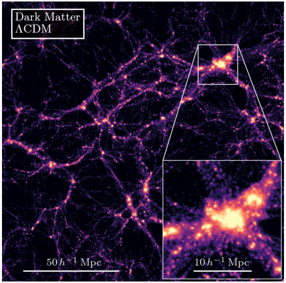
Ferlito, SV, Mota, Baldi, arXiv:2201.04528 (to appear in MNRAS)

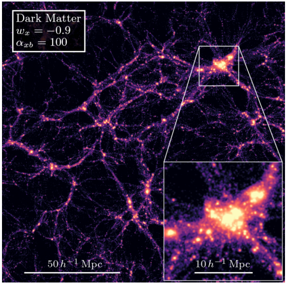
N-body simulations of DE-baryon scattering

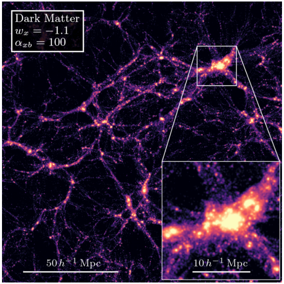
Simulation snapshots:

- $\sigma = 100\sigma_T$
- $w = -0.9, -1, -1.1$

Ferlito, SV, Mota, Baldi, arXiv:2201.04528 (to appear in MNRAS)



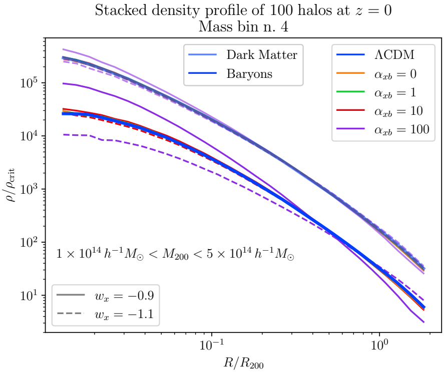




N-body simulations of DE-baryon scattering

Other observables:

- (Cumulative) halo mass function
- (Stacked) halo density profiles
- Baryon fraction profiles
- Future work: Bullet-like systems, higher-order correlators, galaxy bias



Ferlito, SV, Mota, Baldi, arXiv:2201.04528 (to appear in MNRAS)

Baryon profiles most promising observable to probe DE-baryon scattering

Recap

Direct detection of dark energy

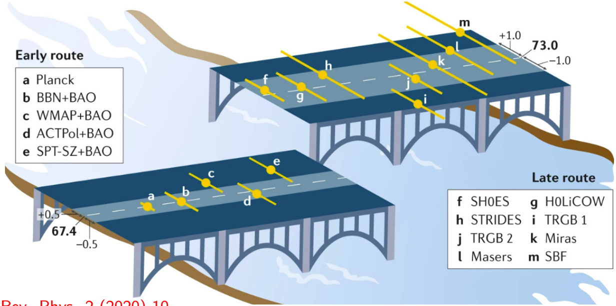
- Potentially lots of unharvested potential for direct detection of dark energy in dark matter direct detection experiments
- Room for large dark energy-baryons interactions in cosmology...
- ...possibly tightly constrained by (non-linear) LSS clustering and other astrophysical observations!

Where else might we learn something about dark energy (at early and late times)?

Perhaps from the Hubble tension!

Part II:
consistency tests of ΛCDM and implications for (early and late) DE

Viewing the Hubble tension ocean with different eyeglasses



Credits: Riess, Nat. Rev. Phys. 2 (2020) 10

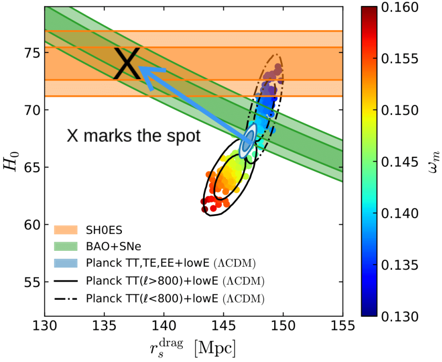
Why does ΛCDM fit data so well? Do we really need new physics? If so, at what time(s), and with what ingredients?

Consistency tests of ΛCDM

The Hubble tension and new physics

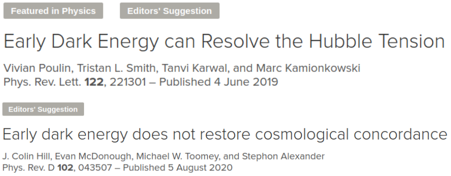
Hubble tension *appears* to call for (substantial) early-time new physics...

Increasing $H(z)$ just prior to $z_\star$: "least unlikely" proposal?



Credits: Knox & Millea, PRD 101 (2020) 043533

Example: early dark energy (some debate as to how much it works)


Featured in Physics | Editors' Suggestion

Early Dark Energy can Resolve the Hubble Tension

Vivian Poulin, Tristan L. Smith, Tanvi Karwal, and Marc Kamionkowski
Phys. Rev. Lett. **122**, 221301 – Published 4 June 2019

Editors' Suggestion

Early dark energy does not restore cosmological concordance

J. Colin Hill, Evan McDonough, Michael W. Toomey, and Stephon Alexander
Phys. Rev. D **102**, 043507 – Published 5 August 2020

Need $\approx$ 12% (!!!) EDE around $z_{\rm eq}$ $\Downarrow$

Why is there no clear sign of new physics in CMB data alone?

Caveat: true prior to ACT DR4 and SPT-3G?

Early-time consistency tests of ΛCDM

PHYSICAL REVIEW D **104,** 063524 (2021)

Consistency tests of ΛCDM from the early integrated Sachs-Wolfe effect: Implications for early-time new physics and the Hubble tension

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(Received 15 June 2021; accepted 22 July 2021; published 15 September 2021)

No clear sign of early-time new physics in CMB data alone

↓

Why does ΛCDM fit CMB data so well?

↓

(Early-time) Consistency tests of ΛCDM

The early ISW (eISW) effect

Around recombination: Universe not fully matter dominated $\Longrightarrow$ residual decay of gravitational potentials $\Longrightarrow$ eISW effect sources anisotropies

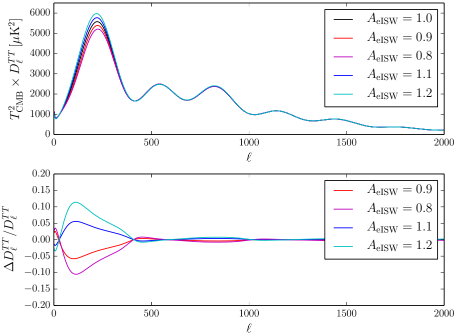
$$\Theta = \int_0^{\eta_0} d\eta \left[\underbrace{\propto g(\Theta_0 + \Psi)}_{\text{Sachs-Wolfe}} + \underbrace{\propto g v_b \frac{d}{d\eta}}_{\text{Doppler}} + \underbrace{\propto e^{-\tau}(\dot{\Psi} - \dot{\Phi})}_{\text{ISW}} + \underbrace{\propto (g\Pi + [g\ddot{\Pi}])}_{\text{Polarization}} \right] j_\ell(k\Delta\eta)$$

$$\Theta_\ell^{\rm ISW}(k) = \underbrace{\int_0^{\eta_m} d\eta \, e^{-\tau} \left(\dot{\Psi} - \dot{\Phi}\right) j_\ell(k\Delta\eta)}_{\text{early ISW}} + \underbrace{\int_{\eta_m}^{\eta_0} d\eta \, e^{-\tau} \left(\dot{\Psi} - \dot{\Phi}\right) j_\ell(k\Delta\eta)}_{\text{late ISW}}$$

(A substantial amount of) New physics increasing $H(z)$ around $z_{\rm eq}/z_\star$ *should* leave an imprint on the eISW effect!

eISW consistency test

$$\Theta_\ell^{\rm eISW}(k) = A_{\rm eISW} \int_0^{\eta_m} d\eta\, e^{-\tau} \left(\dot\Psi - \dot\Phi\right) j_\ell(k\Delta\eta)$$

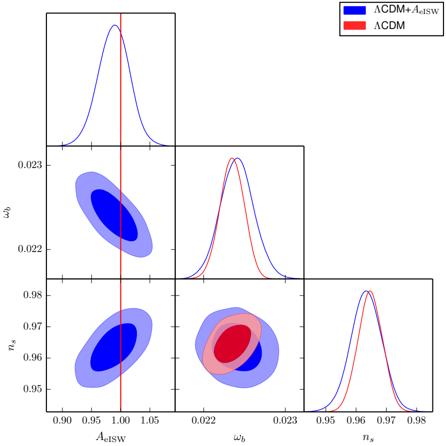


SV, PRD 104 (2021) 063524

Consistency check: within ΛCDM, data consistent with $A_{\rm eISW} = 1$?

eISW consistency test

Is the data consistent with $A_{\rm eISW} = 1$? (7-parameter ΛCDM+$A_{\rm eISW}$)



SV, PRD 104 (2021) 063524

Yes!

Parameter	*Planck* ΛCDM	*Planck* ΛCDM+$A_{\rm eISW}$
$100\omega_b$	2.235 ± 0.015	2.241 ± 0.020
ω_c	0.1202 ± 0.0013	0.1203 ± 0.0014
θ_s	1.0409 ± 0.0003	1.0409 ± 0.0003
τ	0.0544 ± 0.0078	0.0541 ± 0.0078
$\ln(10^{10}A_s)$	3.045 ± 0.016	3.046 ± 0.016
n_s	0.965 ± 0.004	0.963 ± 0.005
$A_{\rm eISW}$	1.0	0.988 ± 0.027
H_0 [km/s/Mpc]	67.26 ± 0.57	67.28 ± 0.62
Ω_m	0.317 ± 0.008	0.317 ± 0.009

SV, PRD 104 (2021) 063524

Other parameter constraints very stable, no more than $\approx 0.3\sigma$ shifts

Implications for early-time new physics: EDE case study

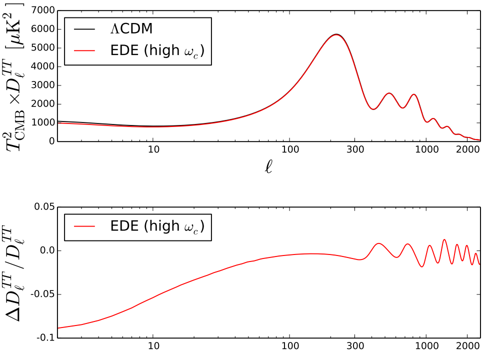
High H_0 EDE fit to CMB at the cost of increase in ω_c $\rightarrow$ worsens tension with WL/LSS data? Hill *et al.*, PRD 102 (2020) 043507; Ivanov *et al.*, PRD 102 (2020) 103502; D'Amico *et al.*, JCAP 2105 (2021) 072; see partial rebuttals in: Murgia *et al.*, PRD 103 (2021) 063502; Smith *et al.*, PRD 103 (2021) 123542

Editors' Suggestion

Early dark energy does not restore cosmological concordance

J. Colin Hill, Evan McDonough, Michael W. Toomey, and Stephon Alexander
Phys. Rev. D **102**, 043507 – Published 5 August 2020

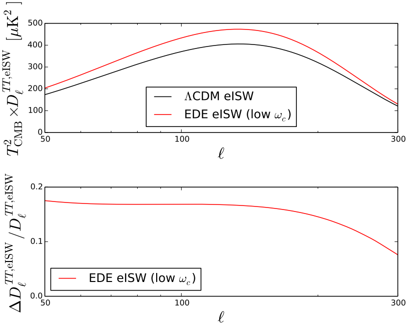
Parameter	ΛCDM	EDE (high ω_c)	EDE (low ω_c)
$100\omega_b$	2.253	2.253	2.253
ω_c	0.1177	0.1322	0.1177
H_0 [km/s/Mpc]	68.21	72.19	72.19
τ	0.085	0.072	0.072
$\ln(10^{10}A_s)$	3.0983	3.0978	3.0978
n_s	0.9686	0.9889	0.9889
$f_{\rm EDE}$	–	0.122	0.122
$\log_{10} z_c$	–	3.562	3.562
θ_i	–	2.83	2.83
n	–	3	3



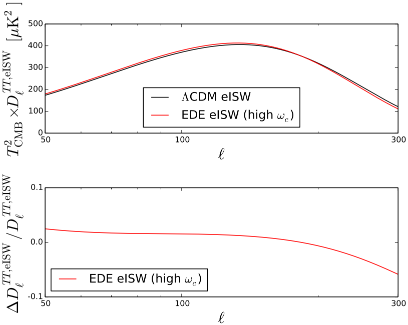
SV, PRD 104 (2021) 063524

Let's extract only the eISW contribution to temperature anisotropies...

Low ω_c



High ω_c



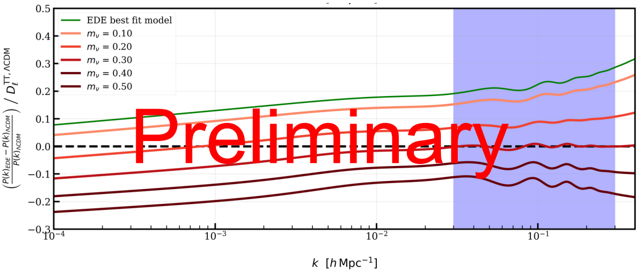
Almost 20% eISW excess!

No more than $\lesssim$ 3-5% eISW excess

Generic to models increasing pre-recombination $H(z)$, not just EDE

Early dark energy problems

Example: neutrino mass (nominally need $M_\nu \sim 0.3\,\mathrm{eV}$ to rescue EDE!)



Reeves, SV, Sherwin, Efstathiou, in preparation. Plot credits: Alex Reeves

Other possible ingredients: decaying DM, DM-dark radiation interactions

Alex Reeves (ETH Zürich)

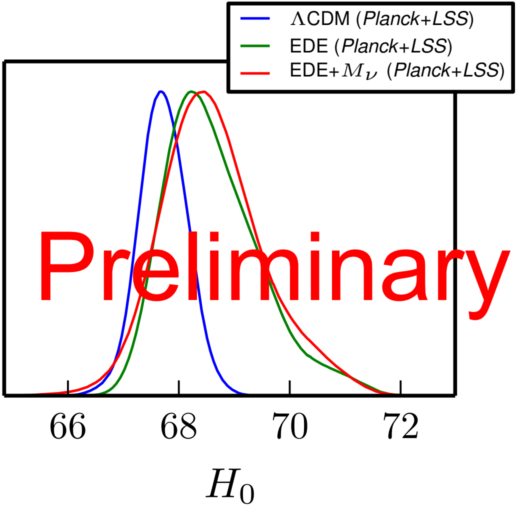
Blake Sherwin (Cambridge)

George Efstathiou (Cambridge)

Early dark energy problems

Massive neutrinos actually turn out not to work:

- Increase in S_8 (worsens S_8 discrepancy)
- M_ν negatively correlated with H_0 for CMB
- Need $M_\nu \sim 0.3\,\mathrm{eV}$, very hard to accommodate in LSS data
- Worsen fit to BAO data
- Maybe EDE not such a bad fit after all (prior volume effects)?



Reeves, SV, Sherwin, Efstathiou, in preparation. Plot credits: Alex Reeves

S_8 discrepancy – something to get excited about?

Monthly Notices of the ROYAL ASTRONOMICAL SOCIETY
MNRAS **505,** 5427–5437 (2021)
Advance Access publication 2021 June 5
https://doi.org/10.1093/mnras/stab1613

Arbitrating the S_8 discrepancy with growth rate measurements from redshift-space distortions

Rafael C. Nunes[1]★ and Sunny Vagnozzi[2]†
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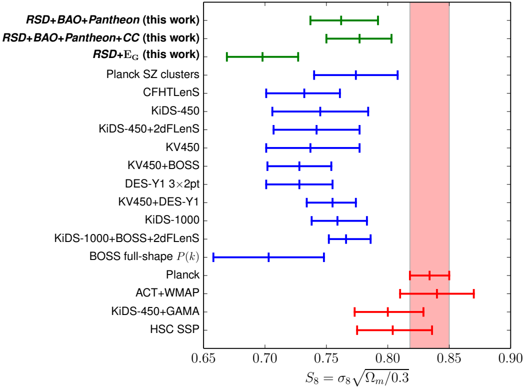
[2]*Kavli Institute for Cosmology (KICC), University of Cambridge, Madingley Road, Cambridge CB3 0HA, UK*



From the growth rate ($f\sigma_8$) point of view, S_8 discrepancy perfectly compatible with a statistical fluctuation!

Rafael Nunes (INPE, Brazil)

Nunes & SV, MNRAS 505 (2021) 5427

Late-time consistency tests of ΛCDM

Is ΛCDM really all there is at late times?

↓

(Try to) Test ΛCDM making no assumptions about early-time physics

↓

Learn something about H_0 in the process?

Old astrophysical objects at high redshift

Historically (1960s-1998) high-z OAO provided the first hints for the existence of dark energy ($\Omega \neq 1$, $\Omega_\Lambda > 0$)

A 3.5-Gyr-old galaxy at redshift 1.55

James Dunlop, John Peacock, Hyron Spinrad, Arjun Dey, Raul Jimenez, Daniel Stern & Rogier Windhorst

Nature **381**, 581–584 (1996) | Cite this article

Conflict over the age of the Universe

M. Bolte & C. J. Hogan

Nature **376**, 399–402 (1995) | Cite this article

The observational case for a low-density Universe with a non-zero cosmological constant

J. P. Ostriker & Paul J. Steinhardt

Nature **377**, 600–602 (1995) | Cite this article

What can OAO do for cosmology in the 2020s?

Cosmology with old astrophysical objects

Can the ages of the oldest inhabitants of the Universe teach us something about the Universe's contents (including DE) and the Hubble tension?

Implications for the Hubble tension from the ages of the oldest astrophysical objects

Sunny Vagnozzi,[1,*] Fabio Pacucci,[2,3,†] and Abraham Loeb[2,3,‡]

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[2]*Center for Astrophysics | Harvard & Smithsonian, Cambridge, MA 02138, USA*
[3]*Black Hole Initiative, Harvard University, Cambridge, MA 02138, USA*

Potentially yes!

Fabio Pacucci (Harvard)

Avi Loeb (Harvard)

Cosmology with old astrophysical objects

$$t_U(z) = \int_z^{\infty} \frac{dz'}{(1+z')H(z')} \propto \frac{1}{H_0}$$

Pros and cons:

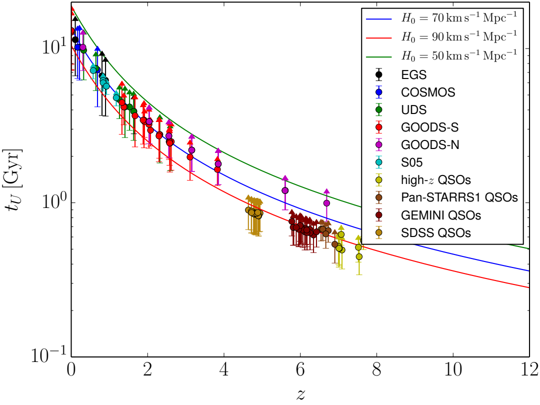
- OAO cannot be older than the Universe $\rightarrow$ **upper limit on** H_0
- $t_U(z)$ integral insensitive to early-time cosmology
- $\implies$ **late-time consistency test for ΛCDM independent of the early-time expansion**!
- Ages of astrophysical objects at $z > 0$ hard to estimate robustly ⚠

Usefulness in relation to the Hubble tension:

- Contradiction between OAO upper limit on H_0 and local H_0 measurements could indicate the need for non-standard late-time ($z \lesssim 10$) physics, or non-standard local physics
- Conclusions completely independent of pre-recombination physics

OAO age-redshift diagram

Age-redshift diagram up to $z \sim 8$

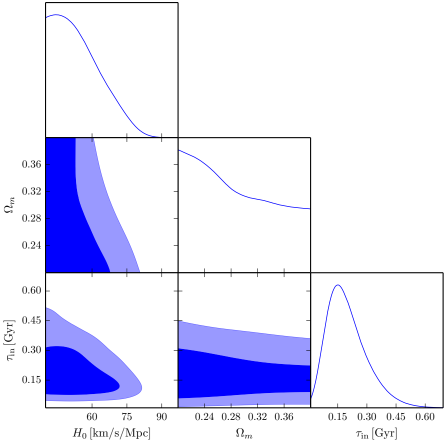


SV *et al.*, arXiv:2105.10421 (submitted to PRL)

Results

Assume ΛCDM at late times, constrain H_0, Ω_m, and incubation time $\tau_{\rm in}$

Prior for $\tau_{\rm in}$ following Jiménez *et al.*, JCAP 1903 (2019) 043; Valcin *et al.*, JCAP 2012 (2020) 022



SV *et al.*, arXiv:2105.10421 (submitted to PRL)

$H_0 < 73.2$ (95% C.L.)

Implications for the Hubble tension

CAVEAT – if the OAO ages are reliable, possible explanations include:

- #1: ΛCDM may not be the end of the story at $z \lesssim 10$
- #2: Nothing wrong with ΛCDM at $z \lesssim 10$, need local new physics...

  Examples: screened 5th forces (Desmond *et al.*, PRD 100 (2019) 043537; Desmond & Sakstein, PRD 102 (2020) 023007), breakdown of FLRW (Krishnan *et al.*, CQG 38 (2021) 184001; arXiv:2106.02532),++
- #3: Just a boring 2σ fluke or systematics?

Is this a hint that pre-recombination new physics alone is not enough to solve the Hubble tension? Krishnan *et al.*, PRD 102 (2020) 103525; Jedamzik *et al.*, Commun. Phys. 4 (2021) 123; Haridasu *et al.*, PRD 103 (2021) 063539; Lin *et al.*, ApJ 920 (2021) 159; Dainotti *et al.*, ApJ 912 (2021) 150

Article | Open Access | Published: 08 June 2021

Why reducing the cosmic sound horizon alone can not fully resolve the Hubble tension

Karsten Jedamzik, Levon Pogosian & Gong-Bo Zhao

Communications Physics **4**, Article number: 123 (2021) | Cite this article

1461 Accesses | **1** Citations | **10** Altmetric | Metrics

Recap

Early-time consistency tests of ΛCDM

- eISW effect sets tight constraints on new pre-recombination physics
- Models which raise pre-recombination $H(z)$ will typically overpredict amplitude of eISW effect
- Example: early dark energy (need additional post-recombination new physics to solve "S_8 tension"?)

Late-time consistency tests of ΛCDM

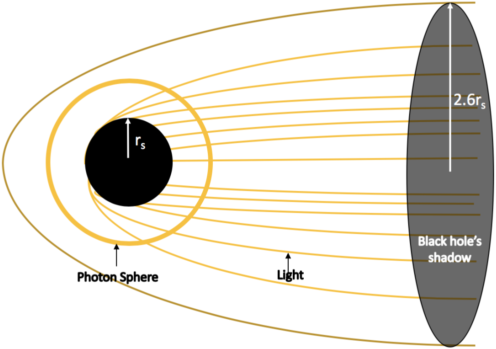
- Ages of old astrophysical objects can set upper limit on H_0
- Late-time consistency test of ΛCDM completely independent of pre-recombination assumptions
- Need new physics at $z \lesssim 10$ or on local scales?

Part III:
new searches for light particles
(dark energy-related or not)

Black hole shadows

For Schwarzschild BH shadow radius $3\sqrt{3}M$

Credits: Event Horizon Telescope collaboration



For advection-dominated hot (geometrically thick optically thin) accretion flow, edge of BH shadow robust to accretion flow details, only influenced by space-time geometry Narayan *et al.*, ApJ Lett. 885 (2019) L33; Bronzwaer & Falcke, arXiv:2108.03966

$\Longrightarrow$ we can use BH shadows to test fundamental physics!

Testing fundamental physics from black hole shadows?

Known information for M87*:

- Diameter of shadow δ, distance to mass ratio D/M
  $\rightarrow d = D\delta/M \sim 11.0 \pm 1.5$
- Deviation from circularity $\Delta C \lesssim 10\%$
- Axis ratio $\Delta y/\Delta x \lesssim 4/3$
- $\epsilon \equiv \Delta Q/Q_{\rm Kerr} \lesssim 4$, $Q_{\rm Kerr} = Ma^2$

Recipe: compute d and ΔC for BHs in your favourite theory, then impose these constraints

Testing the rotational nature of the supermassive object M87* from the circularity and size of its first image

Cosimo Bambi, Katherine Freese, Sunny Vagnozzi, and Luca Visinelli
Phys. Rev. D **100**, 044057 – Published 29 August 2019

Hunting for extra dimensions in the shadow of M87*

Sunny Vagnozzi and Luca Visinelli
Phys. Rev. D **100**, 024020 – Published 12 July 2019

Magnetically charged black holes from non-linear electrodynamics and the Event Horizon Telescope

Alireza Allahyari[1], Mohsen Khodadi[1], Sunny Vagnozzi[2] and David F. Mota[3]
Published 4 February 2020 • © 2020 IOP Publishing Ltd and Sissa Medialab
Journal of Cosmology and Astroparticle Physics, Volume 2020, February 2020
Citation Alireza Allahyari *et al* JCAP02(2020)003

Concerns regarding the use of black hole shadows as standard rulers

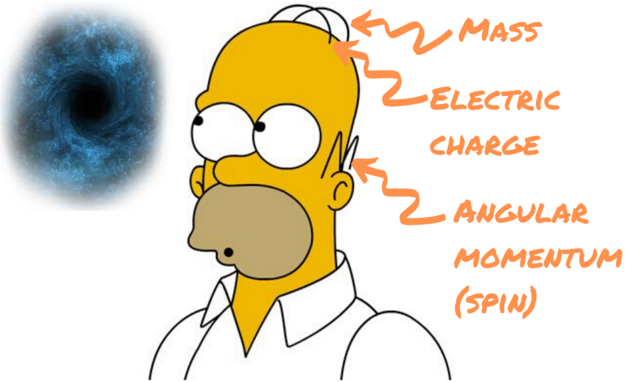
Sunny Vagnozzi[4,1], Cosimo Bambi[2] and Luca Visinelli[3]
Published 25 March 2020 • © 2020 IOP Publishing Ltd
Classical and Quantum Gravity, Volume 37, Number 8
Citation Sunny Vagnozzi *et al* 2020 *Class. Quantum Grav.* **37** 087001

Black holes with scalar hair in light of the Event Horizon Telescope

Mohsen Khodadi[1], Alireza Allahyari[1], Sunny Vagnozzi[2] and David F. Mota[3]
Published 14 September 2020 • © 2020 IOP Publishing Ltd and Sissa Medialab
Journal of Cosmology and Astroparticle Physics, Volume 2020, September 2020
Citation Mohsen Khodadi *et al* JCAP09(2020)026

The no-hair theorem

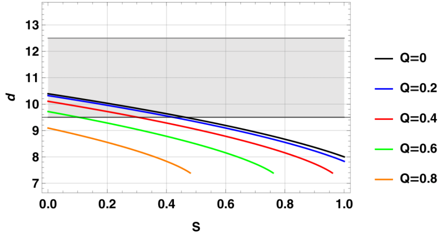
Black holes have at most three hairs ($3 \approx 0$)

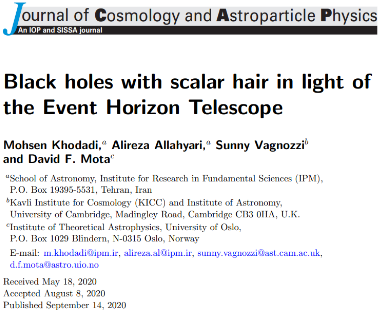


Credits: Medium.com

An example of no-hair theorem violation

$$\mathcal{L} = \mathcal{L}_{\rm EH} + \mathcal{L}_{\rm Maxwell} - \left(\frac{1}{6}\phi^2 R + \partial_\mu \phi \partial^\mu \phi\right)$$




Journal of Cosmology and Astroparticle Physics
An IOP and SISSA journal

Black holes with scalar hair in light of the Event Horizon Telescope

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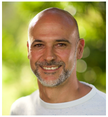
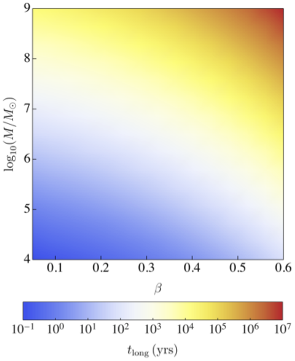


David Mota (Oslo)

Superradiance-induced black hole shadow evolution

Superradiance evolution of black hole shadows revisited

Rittick Roy,[1,*] Sunny Vagnozzi,[2,†] and Luca Visinelli[3,4,‡]
[1] *Center for Field Theory and Particle Physics and Department of Physics, Fudan University, 200438 Shanghai, P. R. China*
[2] *Kavli Institute for Cosmology (KICC) and Institute of Astronomy, University of Cambridge, Madingley Road, Cambridge CB3 0HA, United Kingdom*
[3] *Tsung-Dao Lee Institute (TDLI) and School of Physics and Astronomy, Shanghai Jiao Tong University, 200240 Shanghai, P. R. China*
[4] *INFN, Laboratori Nazionali di Frascati, C.P. 13, 100044 Frascati, Italy*



Roy, SV, Visinelli, arXiv:2112.06932 (submitted to PRD)

- Consistent modelling of superradiant instability in the presence of gas accretion and GW emission
- Evolution in size of SgrA* shadow $\Delta\theta \sim \mathcal{O}(1)\mu\mathrm{as}$ due to superradiance potentially observable on human timescales $[\mathcal{O}(10)\,\mathrm{yr}]$

Rittick Roy (Fudan)

Luca Visinelli (Shanghai)

Precession of planetary objects and new light particles

Asteroid astrometry as a fifth-force and ultralight dark sector probe

Yu-Dai Tsai,[1,2,*] Youjia Wu,[3,†] Sunny Vagnozzi,[4,‡] and Luca Visinelli[5,6,§]

[1]*Fermi National Accelerator Laboratory (Fermilab), Batavia, IL 60510, USA*
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[6]*Tsung-Dao Lee Institute (TDLI), Shanghai Jiao Tong University, 200240 Shanghai, China*

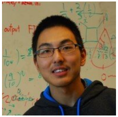
Yukawa potential from new light particle, e.g. new scalar or vector mediator from gauged $U(1)'$ sector $[U(1)_B,\ U(1)_{B-L},\ L_\mu - L_{e,\tau},\ldots]$:

$$V(r) = \tilde{\alpha}\frac{GM_\odot M_*}{r}\exp\left[-\frac{r}{\lambda}\right] \to \frac{g^2}{4\pi}\frac{Q_\odot\, Q_*}{r}\exp\left[-\frac{mc^2}{\hbar c}r\right]\,,$$

Several GR calculations later, in the light mediator limit $(m \ll \hbar/ac)$...

$$|\Delta\varphi| \simeq \frac{2\pi}{1+\frac{g^2}{4\pi G m_p^2}}\frac{g^2}{4\pi G m_p^2}\left(\frac{amc}{\hbar}\right)^2(1-\mathrm{e})$$

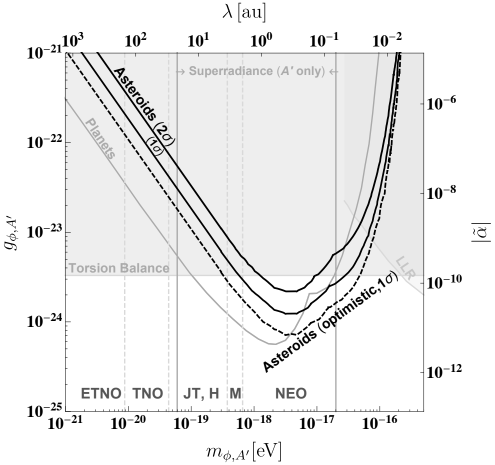

Yu-Dai Tsai (Fermilab → Irvine)


Youjia Wu (Michigan)


Luca Visinelli (Shanghai)

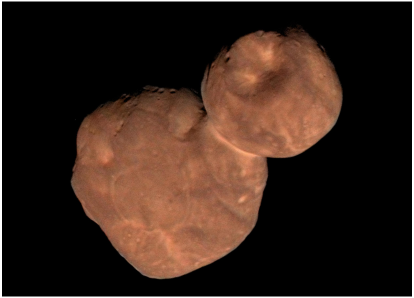
Precession of planetary objects and new light particles

Results from planets and 9 well-tracked (i.e. dangerous) asteroids



Tsai, Wu, SV, Visinelli, arXiv:2107.04038 (submitted to Nat. Astron.)

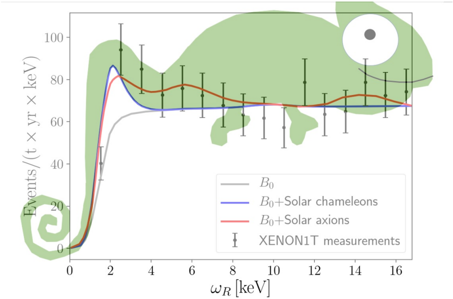
Asteroid 66391 Moshup
($\sim$ 1.3 km in diameter)



Credits: Wikiwand

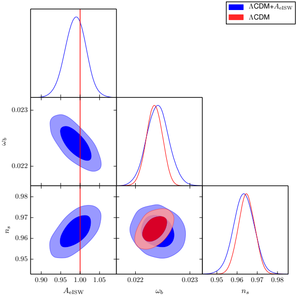
Conclusions

Direct detection of dark energy: lots of unharvested potential in dark matter direct detection experiments



SV *et al.*, PRD 104 (2021) 063023

Consistency tests of ΛCDM: pre-recombination new physics tightly constrained by eISW effect



SV, PRD 104 (2021) 063524

Much to be learned about dark energy beyond "standard" cosmological searches for its gravitational interactions