Seven hints that early-time new physics alone is not sufficient to solve the Hubble tension

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University of Zurich, 9 April 2024

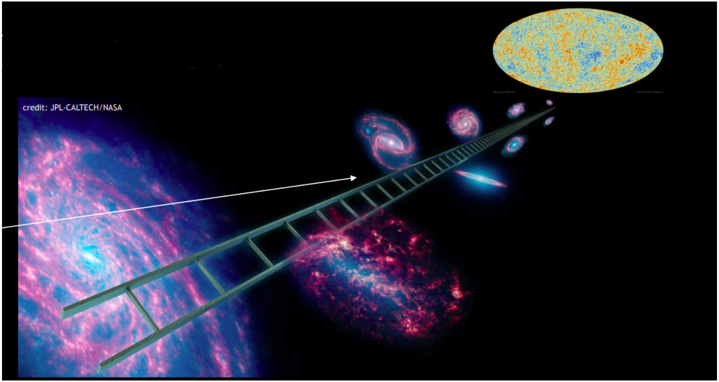
The Hubble constant

H_0: current rate of expansion of the Universe

Why care about H_0?

- Allan Sandage, 1970: "*Cosmology can be described as the search for two numbers: the current rate of expansion [H_0] and the deceleration of the expansion [q_0]*"
- Adam Riess, 2019: "*H_0 is the ultimate end-to-end test for ΛCDM*"

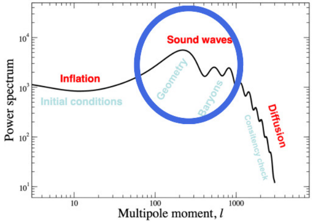
H_0 as an end-to-end test

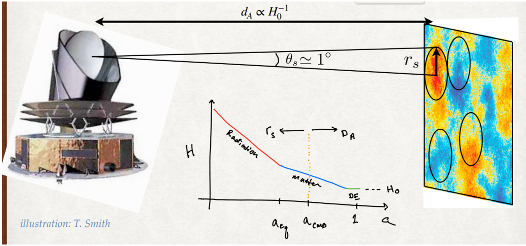


Credits: JPL-Caltech/NASA and Dillon Brout

How to measure H_0? Always a good idea in cosmology: measure distances (or rather, infer distances from angles and fluxes using standard rulers and standard candles!)

The CMB as a (self-calibrated) standard ruler

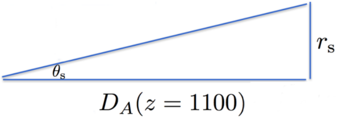




Credits: Planck collaboration and Silvia Galli (left); Tristan Smith and Vivian Poulin (right)

$$\theta_s = \frac{r_s}{d_A(z_\star)} = 0.010411 \pm \underline{\underline{0.000003}} \quad (!!!)$$

Note: θ_s measured exquisitely, but r_s and d_A are model-dependent!



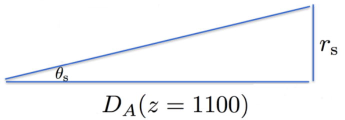
Credits: Silvia Galli

Steps to apply the CMB ruler

Within ΛCDM:

$$\theta_s = \frac{r_s}{d_A(z_\star)}\,, \qquad r_s \simeq \int_{z_\star}^{\infty} dz\, \frac{c_s(z,\omega_b,\omega_r)}{\sqrt{(\omega_c+\omega_b)(1+z)^3+\omega_r(1+z)^4}}$$

- ω_r: exquisitely measured from T_{CMB} (e.g. COBE)
- $c_s(z) = (1+3\rho_b/4\rho_\gamma)^{-1}$
- ω_b: infer from relative height of odd and even peaks, further improvement from damping tail
- ω_c: infer from early ISW effect (first peak height), potential envelope, further improvement from lensing-induced peak smoothing



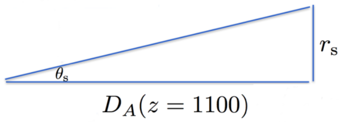
Credits: Silvia Galli

Steps to apply the CMB ruler

Within ΛCDM:

$$\theta_s = \frac{r_s}{d_A(z_\star)}\,, \qquad d_A(z_\star) \simeq 3\int_0^{z_\star} dz\,\frac{1}{\sqrt{\omega_\Lambda + \omega_m(1+z)^3 + \omega_r(z)}}\,\text{Gpc}$$

- $\omega_r(z)$: already known as before
- $\omega_m = \omega_c + \omega_b$: both terms already known as before
- θ_s: inferred from peak spacing, $\theta_s \simeq \pi/\Delta\ell = \pi/(\ell_{p+1} - \ell_p)$
- ω_Λ: only remaining free parameter, to fix from $d_A(z_\star) = r_s\Delta\ell/\pi$
- Once ω_Λ is known, the whole evolution of $H(z)$ is known, including $H(z=0) = H_0$!



Credits: Silvia Galli

Applying the ruler

Units of H_0 always implicitly km/s/Mpc from now

$$H_0 = 67.27 \pm 0.60$$
$$(Planck\ 2018\ TTTEEE{+}lowE)$$

Planck collaboration, A&A 641 (2020) A6

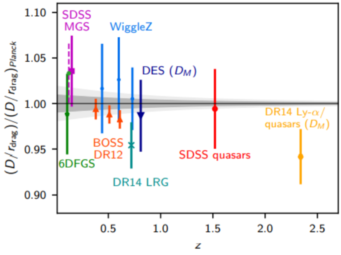
$$H_0 = 67.9 \pm 1.5$$
$$(ACT\ DR4)$$

ACT collaboration, JCAP 2012 (2020) 047

Late-time guard rails

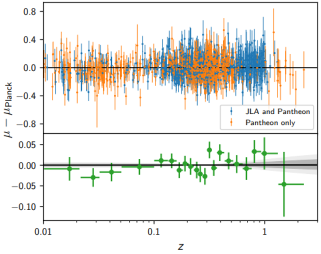
It is important to "stabilize" CMB-only constraints with late-time datasets, *especially when going beyond ΛCDM at late times*!

BAO



Planck collaboration, A&A 641 (2020) A6

Cosmological/high-z SNeIa



Planck collaboration, A&A 641 (2020) A6

These are in *very good* agreement with the expansion history inferred from *Planck* within ΛCDM (so in ΛCDM mostly a consistency check)!

Combining CMB and late-time guard rails

Combination consistent with CMB-only value of H_0 within ΛCDM, important sanity check!

$$H_0 = 67.72 \pm 0.40 \; (CMB+BAO+uncalibrated\ SNeIa)$$

Planck collaboration, A&A 641 (2020) A6

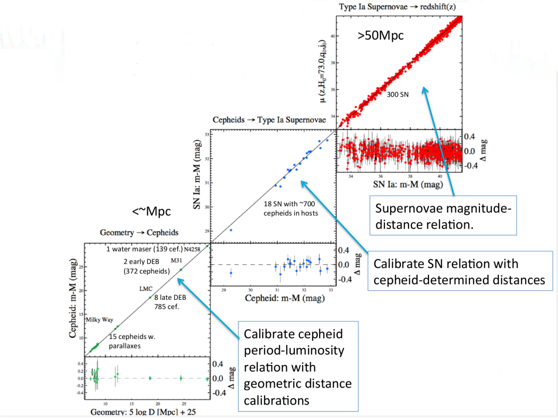
Fresh results from DESI (as of last week):

$$H_0 = 67.97 \pm 0.38 \; (CMB+DESI\ BAO)$$

DESI collaboration, arXiv:2404.03002

Calibrating the local distance ladder with Cepheids

Best known 3-rung distance ladder: Cepheid-calibrated SNeIa



Credits: adapted from Adam Riess and Silvia Galli

Applying the ladder

SH0ES analysis: 75 MW Cepheids with *Gaia* EDR3 parallaxes (plus other geometric distances), $>$90 Cepheids, 42 calibrator SNeIa in 37 SNeIa+Cepheid hosts, 277 SNeIa in $0.0233 < z < 0.15$
$\implies$ 1.4% measurement of H_0!

$$H_0 = 73.04 \pm 1.04$$
(Cepheid-calibrated SNeIa, R22)

Riess *et al.*, ApJ Lett. 934 (2022) L7

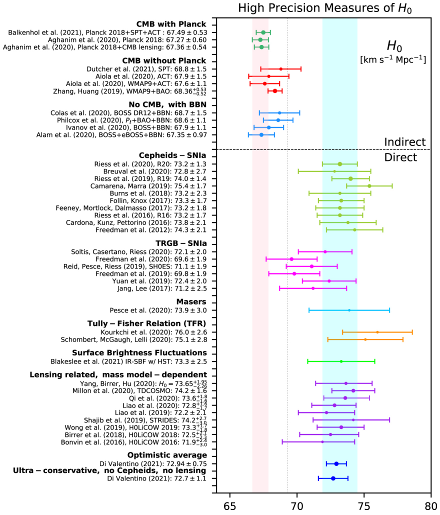
Notes:

- need intermediate rung as SNeIa are rare events, not enough of them in the local Universe for direct parallax calibration
- Cepheids are standard candles through period-luminosity relation

The trouble

Overall trend:

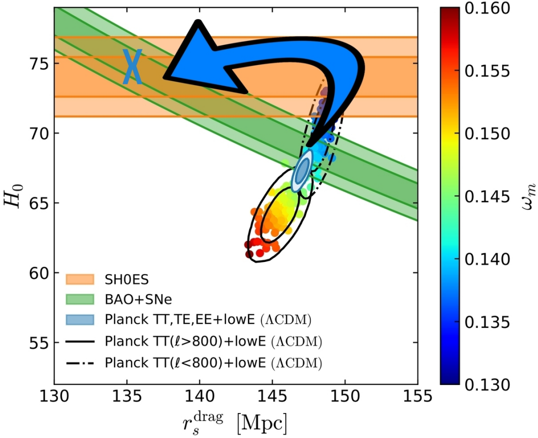
- "early-time" model-dependent measurements prefer low H_0
- "late-time" direct measurements prefer high H_0



Review by Di Valentino *et al.*, CQG 38 (2021) 153001

Hubble tension "no-go theorem"

Solving the tension while providing a good fit to BAO data and Hubble flow SNeIa data seems to require lowering r_s by $\approx 7\%$



Knox & Millea, PRD 101 (2020) 043533

This would seem to require early-time (pre-recombination) new physics!

Hubble tension "no-go theorem"?

...yet, we still haven't been able to construct a model truly fixing H_0 (empirically, early-Universe new physics only seems to get to $H_0 \sim 70$ – *with Planck CMB data and without including local H_0 priors*)

Is early-time new physics the end of the story?

My sociological worry: "*the Hubble tension calls for early-time new physics*" may have been uncritically elevated to the mantra "*the Hubble tension calls* **exclusively** *for early-time new physics*"

Seven hints

- 𝒜ges of the oldest astrophysical objects
- ℬaryon Acoustic Oscillations r_d-H_0 degeneracy slope
- 𝒞osmic chronometers
- 𝒟escending trends observed in a wide range of low-z datasets
- ℰarly integrated Sachs-Wolfe effect and its restrictions on early-time new physics
- ℱractional matter density (Ω_m) constraints from uncalibrated cosmic standards
- 𝒢alaxy power spectrum r_d- and $k_{\rm eq}$-based determinations of H_0

Why seven? (*Why not?*) Miller's law – see Miller, Psychol. Rev. 63 (1956) 81

universe MDPI

Opinion

Seven Hints that Early-Time New Physics Alone Is Not Sufficient to Solve the Hubble Tension

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SV, Universe 9 (2023) 393

Seven hints

a) Just reducing the sound horizon will introduce other problems:

- $\mathcal{B}$aryon Acoustic Oscillations r_d-H_0 degeneracy slope

b) Early-time guard rails – introducing pre-recombination new physics and maintaining the level of early-time consistency of ΛCDM is difficult:

- $\mathcal{E}$arly integrated Sachs-Wolfe effect and its restrictions on early-time new physics
- $\mathcal{G}$alaxy power spectrum r_d- and $k_{\rm eq}$-based determinations of H_0

c) Analyses more-or-less independent of pre-recombination physics – some residual amount of post-recombination physics seems to be required:

- $\mathcal{F}$ractional matter density (Ω_m) constraints from uncalibrated cosmic standards
- $\mathcal{C}$osmic chronometers
- $\mathcal{A}$ges of the oldest astrophysical objects
- $\mathcal{D}$escending trends observed in a wide range of low-z datasets

Just reducing the sound horizon will introduce other problems

Hint 1: BAO r_d-H_0 degeneracy slope

CMB and BAO constrain respectively:

$$\theta_\star \equiv \frac{r_\star}{D(z_\star)}\,, \qquad \theta_d(z_{\rm obs}) \equiv \frac{r_d}{D(z_{\rm obs})}$$

Two sound horizons closely related:

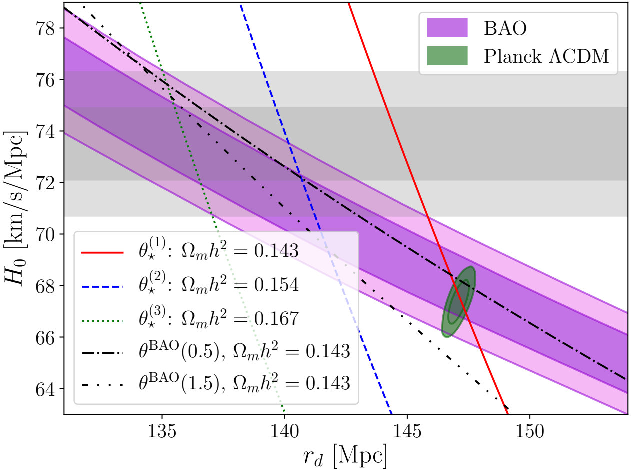
$$r_d \approx 1.0184 r_\star$$

Given ω_m, imposing $\theta_\star = \text{const}$ and $\theta_d(z_{\rm obs}) = \text{const}$ defines degeneracy line in r_d-H_0 plane with very different slopes for CMB and BAO (steeper for CMB, because $z_\star \gg z_{\rm obs}$)

Q: what happens if H_0 is raised while *only* lowering r_d...?

Hint 1: BAO r_d-H_0 degeneracy slope

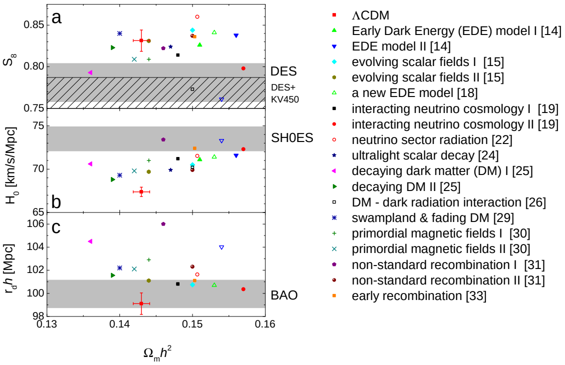
A: quickly run into trouble with BAO and/or WL data if ω_m is unchanged, but even changing ω_m cannot bring agreement with both!



Jedamzik, Pogosian & Zhao, Commun. Phys. 4 (2021) 123

Hint 1: BAO r_d-H_0 degeneracy slope

Lower $\omega_m \implies$ tension with BAO data
Higher $\omega_m \implies$ tension with WL data (worsen S_8 tension)



Jedamzik, Pogosian & Zhao, Commun. Phys. 4 (2021) 123

New physics which *only* reduces r_s is not enough!

Early-time guard rails

Hint 2: Early ISW effect

Around recombination: Universe not fully matter dominated $\implies$ residual decay of gravitational potentials $\implies$ eISW effect sources anisotropies

$$\Theta = \int_0^{\eta_0} d\eta \left[\underbrace{\propto g(\Theta_0 + \Psi)}_{\text{Sachs-Wolfe}} + \underbrace{\propto g v_b \frac{d}{d\eta}}_{\text{Doppler}} + \underbrace{\propto e^{-\tau}(\dot{\Psi} - \dot{\Phi})}_{\text{ISW}} + \underbrace{\propto (g\Pi + [g\ddot{\Pi}])}_{\text{Polarization}} \right] j_\ell(k\Delta\eta)$$

$$\Theta_\ell^{\rm ISW}(k) = \underbrace{\int_0^{\eta_m} d\eta\, e^{-\tau} \left(\dot{\Psi} - \dot{\Phi}\right) j_\ell(k\Delta\eta)}_{\text{early ISW}} + \underbrace{\int_{\eta_m}^{\eta_0} d\eta\, e^{-\tau} \left(\dot{\Psi} - \dot{\Phi}\right) j_\ell(k\Delta\eta)}_{\text{late ISW}}$$

(A substantial amount of) New physics increasing $H(z)$ around $z_{\rm eq}/z_\star$ *should* leave an imprint on the eISW effect!

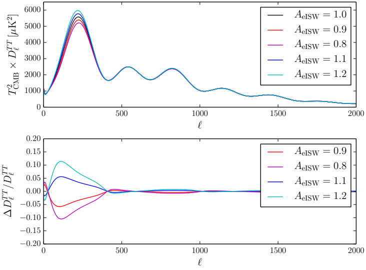
Why is there no clear sign of early-time new physics in CMB data alone?

Hint 2: Early ISW effect

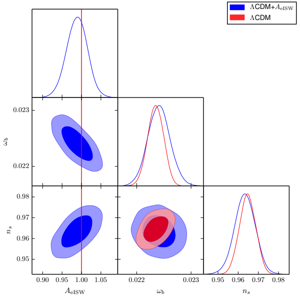
$$\Theta_\ell^{\text{eISW}}(k) = A_{\text{eISW}} \int_0^{\eta_m} d\eta \, e^{-\tau} \left(\dot{\Psi} - \dot{\Phi}\right) j_\ell(k\Delta\eta)$$

Consistency check: within ΛCDM, data consistent with $A_{eISW} = 1$?

Yes! $A_{\text{eISW}} = 0.988 \pm 0.027$ (other parameters stable to within $\lesssim 0.3\sigma$)



SV, PRD 104 (2021) 063524



SV, PRD 104 (2021) 063524

Hint 2: Early ISW effect (EDE application)

High H_0 EDE fit to CMB requires increased ω_c $\rightarrow$ worsens S_8 tension?

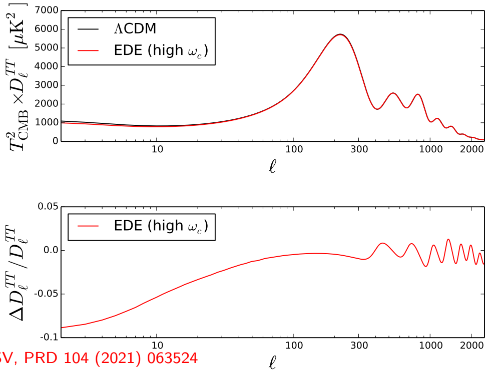
Hill *et al.*, PRD 102 (2020) 043507; Ivanov *et al.*, PRD 102 (2020) 103502; D'Amico *et al.*, JCAP 2105 (2021) 072; see partial rebuttals in: Murgia *et al.*, PRD 103 (2021) 063502; Smith *et al.*, PRD 103 (2021) 123542

Editors' Suggestion

Early dark energy does not restore cosmological concordance

J. Colin Hill, Evan McDonough, Michael W. Toomey, and Stephon Alexander
Phys. Rev. D **102**, 043507 – Published 5 August 2020

Parameter	ΛCDM	EDE (high ω_c)	EDE (low ω_c)
$100\omega_b$	2.253	2.253	2.253
ω_c	0.1177	0.1322	0.1177
H_0 [km/s/Mpc]	68.21	72.19	72.19
τ	0.085	0.072	0.072
$\ln(10^{10}A_s)$	3.0983	3.0978	3.0978
n_s	0.9686	0.9889	0.9889
$f_{\rm EDE}$	–	0.122	0.122
$\log_{10} z_c$	–	3.562	3.562
θ_i	–	2.83	2.83
n	–	3	3

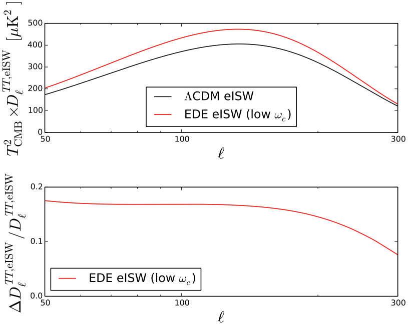


SV, PRD 104 (2021) 063524

Hint 2: Early ISW effect (EDE application)

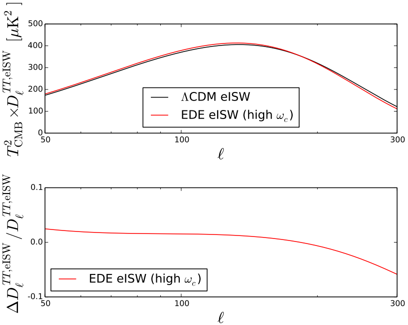
Let's extract only eISW contribution to temperature anisotropies...

Low ω_c



SV, PRD 104 (2021) 063524

High ω_c



SV, PRD 104 (2021) 063524

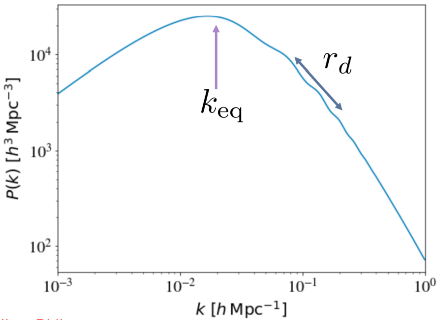
Almost 20% eISW excess!

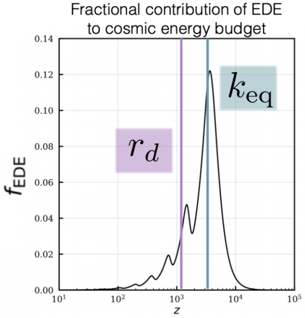
No more than $\lesssim$ 3-5% eISW excess

Problem generic to models increasing pre-recombination $H(z)$

Hint 3: r_s- and k_{eq}-based constraints on H_0 from $P(k)$

Two scales in $P(k)$, both standard rulers





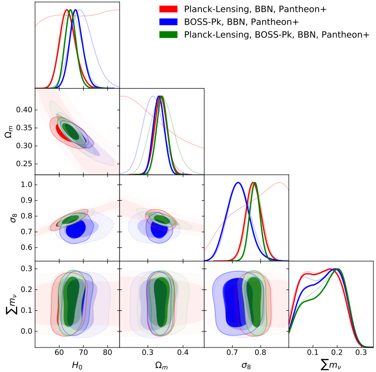
Credits: Oliver Philcox

- $k_{eq} = \sqrt{2\Omega_m H_0 z_{eq}}$ (if no extra components with significant pressure support) sets peak and overall shape ($z_{eq} \approx 3500$)
- r_d sets BAO frequency ($z_\star \approx 1100$)

Both can be used to infer H_0: in the presence of a substantial amount of early-time new physics, no reason two values should agree!

Hint 3: r_s- and $k_{\rm eq}$-based constraints on H_0 from $P(k)$

Can analyze $P(k)$ data removing (most) r_d information (effectively marginalizing over r_d), similarly CMB lensing also sensitive to $k_{\rm eq}$



Philcox *et al.*, PRD 106 (2022) 063530

$H_0 = 64.8^{+2.2}_{-2.5}$ (only $k_{\rm eq}$ info): agrees with ΛCDM r_d-based value of H_0, disfavors significant amount of early-time new physics?

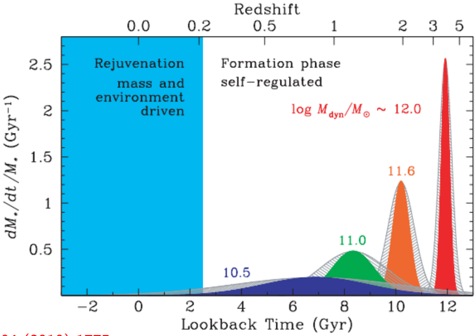
Analyses more-or-less independent of pre-recombination physics

Hint 5: Cosmic chronometers

Take two ensembles of galaxies that formed around the same time and are separated by a small redshift interval Δz around z_{eff}: Jiménez & Loeb, ApJ 573 (2002) 37

$$\frac{dt}{dz} = -\frac{1}{(1+z)H(z)} \implies H(z_{\text{eff}}) = -\frac{1}{1+z_{\text{eff}}}\frac{\Delta z}{\Delta t}$$

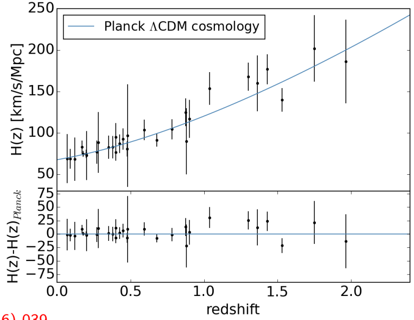
Use massive, early-time, passively-evolving galaxies (evolving on a much longer timescale than their age differences)



Thomas *et al.*, MNRAS 404 (2010) 1775

Hint 5: Cosmic chronometers

- CCs are completely (cosmological) model-independent
- CCs can be used to infer *cosmological*/*non-local* value of H_0
- Analyzing CC requires no assumptions on early-Universe physics
- Contradiction between CCs value of H_0 (assuming ΛCDM) and local H_0 measurements could indicate the need for non-standard late-time ($z \lesssim 2$) physics beyond ΛCDM, or non-standard local physics

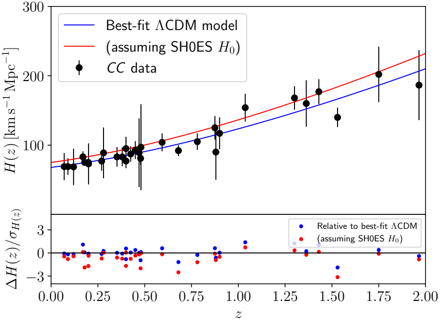


Moresco *et al.*, JCAP 1612 (2016) 039

Hint 5: Cosmic chronometers

Early-time-independent consistency test of ΛCDM: assuming ΛCDM holds at late times, from CC alone infer $H_0 = 67.5 \pm 3.0$ (note: no systematics!)

- Central value in excellent agreement with *Planck*
- Almost 2σ "tension" with local Cepheid-calibrated SNeIa H_0
- Preference for low H_0 not driven by any specific datapoint
- If uncertainties decrease and central value doesn't move, will need new late-time ($z \lesssim 2$) physics and/or new local physics



Hint 6: Ages of the oldest astrophysical objects

Historically (1960s-1998) high-z OAOs provided the first hints for the existence of dark energy ($\Omega \neq 1$, $\Omega_\Lambda > 0$)

A 3.5-Gyr-old galaxy at redshift 1.55

James Dunlop, John Peacock, Hyron Spinrad, Arjun Dey, Raul Jimenez, Daniel Stern & Rogier Windhorst

Nature **381**, 581–584 (1996) | Cite this article

The observational case for a low-density Universe with a non-zero cosmological constant

J. P. Ostriker & Paul J. Steinhardt

Nature **377**, 600–602 (1995) | Cite this article

What can OAOs do for cosmology in the 2020s?

Hint 6: Ages of the oldest astrophysical objects

$$t_U(z) = \int_z^{\infty} \frac{dz'}{(1+z')H(z')} \propto \frac{1}{H_0}$$

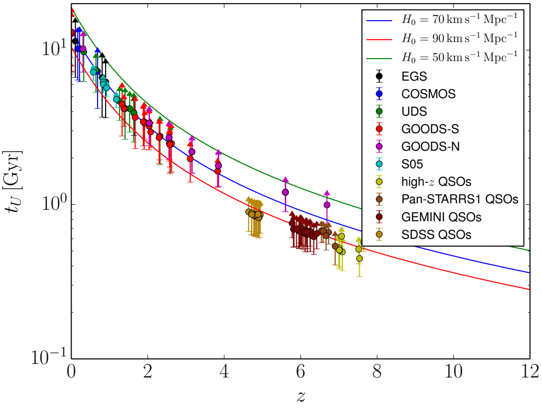
- OAOs cannot be older than the Universe → **upper limit on** H_0
- $t_U(z)$ integral insensitive to early-time cosmology
- → **late-time ΛCDM consistency test independent of early times**!
- Ages of astrophysical objects at $z > 0$ hard to estimate robustly ⚠

Usefulness in relation to the H_0 tension:

- Contradiction between OAOs upper limit on H_0 and local H_0 measurements could indicate the need for non-standard late-time ($z \lesssim 10$) physics, or non-standard local physics
- Conclusions completely independent of pre-recombination physics

Hint 6: Ages of the oldest astrophysical objects

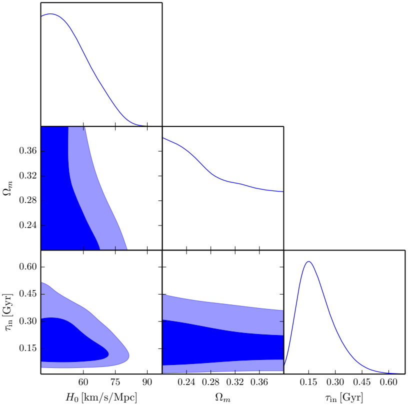
Age-redshift diagram up to $z \sim 8$



SV, Pacucci & Loeb, JHEAp 36 (2022) 27

Hint 6: Ages of the oldest astrophysical objects

Assume ΛCDM at late times, constrain H_0 and Ω_m



SV, Pacucci & Loeb, JHEAp 36 (2022) 27

$H_0 < 73.2$ (95% C.L.)

CAVEAT – If the OAOs ages are reliable, possible explanations are:

1. ΛCDM may not be the end of the story at $z \lesssim 10$
2. Nothing wrong with ΛCDM at $z \lesssim 10$, need new physics on local scales
3. Just a boring 2σ fluke or systematics?

Hint 7: Descending trends

Mathematically speaking, dynamical models (e.g. ΛCDM) break down if values of (constant) fitting parameters pick up time dependence

Integrate 1st Friedmann equation with $w_{\text{eff}}(z)$ prescribed (in FLRW):

$$H_0 = H(z) \exp\left[-\frac{3}{2}\int_0^z dz' \, \frac{1 + w_{\text{eff}}(z')}{1 + z'}\right]$$

$H(z) \sim$ data $\qquad$ $w_{\text{eff}}(z')$: prescribed model
H_0: inferred fitting parameter (here mathematically integration constant)

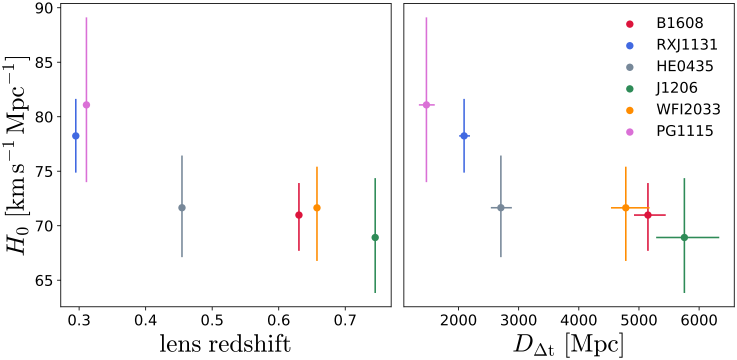
If input $w_{\text{eff}}(z)$ and data "disagree", H_0 picks up z-dependence and "runs" *at all redshifts* Krishnan *et al.*, PRD 103 (2021) 103509

If H_0 tension physical and at least some late-time new physics involved, z-evolution of H_0 at intermediate z ($0 < z < z_\star$) inevitable!

Hint 7: Descending trends

- Has such a z-evolution already been observed in current data?
- Has it been observed in independent datasets with a common trend?
- Are there mundane explanations for its size and direction?

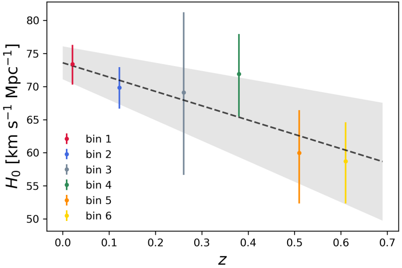
Perhaps most famous example observed in H0LiCOW data ($\sim 2\sigma$)

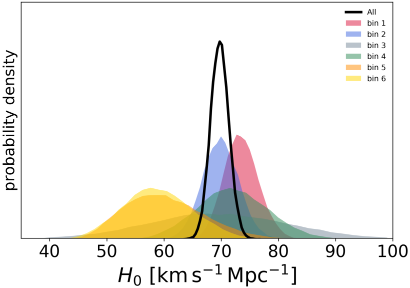


Wong *et al.*, MNRAS 498 (2019) 1420

Hint 7: Descending trends

Combination of (binned) low-z datasets: megamaser distances, CCs, isotropic BAO, *Pantheon* SNeIa (r_d treated as free parameter)





Krishnan *et al.*, PRD 102 (2020) 103525

$\sim 2.1\sigma$ significance, slope consistent with H0LiCOW

Hint 7: Descending trends

Similar trends (descending H_0 and/or increasing Ω_m) observed in many different dataset combinations:

- *Pantheon* SNeIa Dainotti *et al.*, ApJ 912 (2021) 150
- *PantheonPlus+SH0ES* SNeIa Jia, Hu & Wang, A&A 674 (2023) A45
- *PantheonPlus* SNeIa Malekjani *et al.*, arXiv:2301.12725
- *Pantheon* SNeIa Horstmann, Pietschke & Schwarz, A&A 668 (2022) A34
- CC+*Pantheon* SNeIa+QSOs Ó Colgáin *et al.*, arXiv:2206.11447
- QSOs Risaliti & Lusso, Nat. Astron. 3 (2019) 272
- $f\sigma_8$ measurements: S_8 increasing with z Adil *et al.*, MNRAS Lett. 528 (2024) L20
- ...and others!

Question: could this be expected even within ΛCDM? (naïve guess: at high z lose sensitivity to DE, so expect $\Omega_m \uparrow \implies H_0 \downarrow$)

Mock analysis seems to suggest effect is too big and should be seen at higher redshift Ó Colgáin, Sheikh-Jabbari & Solomon, PDU 40 (2023) 101216

Hint 4: Ω_m constraints from uncalibrated cosmic standards

Beneficial to look at joint H_0-Ω_m constraints rather than just projected H_0 constraints Lin, Mack & Hou, ApJL 904 (2020) L22

Can we determine Ω_m:

- At a level competitive with the CMB model-dependent value?
- Free from early-Universe assumptions (as with BAO+SNeIa)?

$\Delta r H_0$ small & insensitive to early-Universe physics Lin, Chen & Mack, ApJ 920 (2021) 159

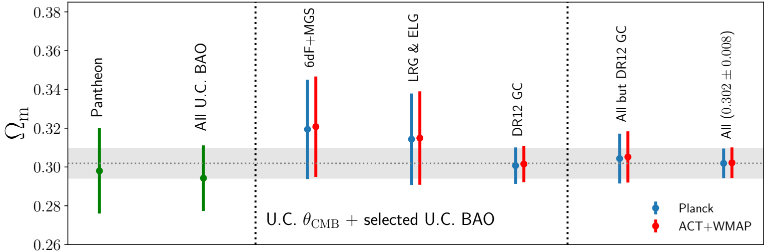
$$\Delta r H_0 \equiv (r_d - r_\star)\, H_0 = \int_{z_d}^{z_\star} dz\, \frac{c_s(z)}{E(z)} \qquad (z_d - z_\star) \sim 30$$

Combine $\theta_\star$ (CMB) and θ_d (BAO) in almost early Universe-independent way, with long lever arm to constrain Ω_m at level competitive with CMB: Early Universe Physics Insensitive Uncalibrated Cosmic Standards (UCS)

Hint 4: Ω_m constraints from uncalibrated cosmic standards

Data: $\theta_\star$ (*Planck*+ACT+WMAP), θ_d (eBOSS), CMB priors on $z_\star$ and Δz_s, BBN prior on $\Omega_b h^2$

Parameters: Ω_m, $\mathcal{M}$, $r_d H_0$, h (weak dependence)



Lin, Chen & Mack, ApJ 920 (2021) 159

Purely geometrical, early Universe-independent value: $\Omega_m = 0.302 \pm 0.008$

For comparison $\Omega_m = 0.310 \pm 0.006$ in ΛCDM using full CMB information

Hint 4: Ω_m constraints from uncalibrated cosmic standards

Constraints not exactly along Ω_m direction, weak Ω_m-h degeneracy

$$\left(\frac{\Omega_m}{0.3}\right)\left(\frac{h}{0.7}\right)^{-0.08} = 1.0060 \pm 0.0258$$

Combine UCS with several early Universe-independent late-time, non-local measurements to infer H_0 in an early Universe-independent way

Methods	H_0 (km s^{-1} Mpc^{-1})		n-σ from R21	
UCS and individual nonlocal observation	Without $\theta_{\rm cmb}$	With $\theta_{\rm cmb}$	Without $\theta_{\rm cmb}$	**With $\theta_{\rm cmb}$**
Cosmic chronometers				
Current public data	69.1 ± 1.7	**68.8 ± 1.6**	1.9σ	**2.1σ**
Extra systematic	69.4 ± 2.3	**69.2 ± 2.1**	1.4σ	**1.6σ**
Extra systematic, conservative	69.3 ± 3.4	**68.9 ± 3.3**	1.1σ	**1.2σ**
γ-ray optical depth	66.2 ± 3.5	66.1 ± 3.4	1.9σ	**2.0σ**
Cosmic age				
$t_{\rm U} = 13.5 \pm 0.27$ Gyr	70.2 ± 1.7	**69.8 ± 1.5**	1.4σ	**1.7σ**
$t_{\rm U} = 13.5 \pm 0.33$ Gyr	70.3 ± 2.1	**69.8 ± 1.9**	1.2σ	**1.5σ**
CMBlens+DES+BBN	68.8 ± 2.4	**68.6 ± 2.0**	1.6σ	**1.9σ**
UCS and joint nonlocal observations[a]				
All nonlocal observations	69.1 ± 1.5	**68.8 ± 1.3**	2.0σ	**2.4σ**
Nonlocal observations without cosmic age	68.3 ± 1.9	**68.1 ± 1.6**	2.1σ	**2.5σ**
Nonlocal observations without LSS	69.1 ± 1.6	**68.8 ± 1.5**	2.0σ	**2.2σ**

Lin, Chen & Mack, ApJ 920 (2021) 159

Residual $\approx 2\sigma$ tension can have nothing to do with early-Universe physics: need late-time new physics and/or local new physics (systematics very unlikely given consistency among independent probes)

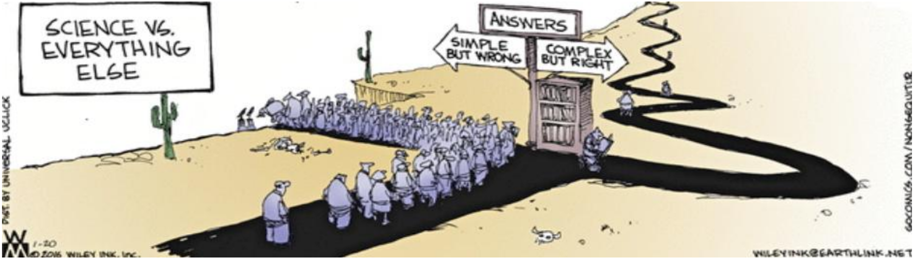
Where to from here? Some scattered thoughts

- Empirically: early-time physics only seems to reach $H_0 \sim 70$ (no external priors)
- Idea: combine early-time and late-time (both non-local) and local new physics?
- Direction of late-time physics: lower $d_A(z)$ at $z > 0$ (phantom/interacting DE?)
- CMB+BAO/SNeIa actually can tolerate w as low as ~ -1.07, H_0 responds as $\Delta H_0 \sim -20(1+w)$, so this can help as much as $\Delta H_0 \sim 1.5$ SV, PRD 102 (2020) 023518
- If there is *also* some local new physics lowering local H_0, maybe don't need non-local H_0 to go all the way up to ~ 74 after all? (two can meet halfway)
- Early-time new physics probably still need to do the lion's share of the job...
- Early+late: can two models decouple, both "push" non-local H_0 up separately, combining their tension-solving virtues "in phase"/"constructively"?

Occam's razor

Objection: wouldn't this violate Occam's razor?

My opinion ↓

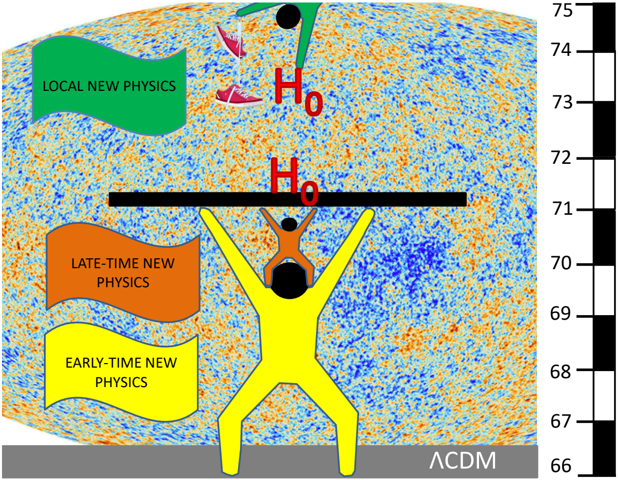


Credits: Wiley Miller

Nature is under no obligation to look simple to us!

Where to from here?

Pictorial representation of what I think could be a promising scenario



Credits: Cristina Ghirardini

Conclusions

Early-time new physics alone cannot solve the Hubble tension – will probably need a combination of early-time and late-time (both non-local) and possibly local new physics

universe MDPI

Opinion

Seven Hints that Early-Time New Physics Alone Is Not Sufficient to Solve the Hubble Tension

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SV, Universe 9 (2023) 393

Backup slides

Phantom dark energy? Really?

The state of the dark energy equation of state circa 2023

Luis A. Escamilla,[1,2,*] William Giarè,[2,†] Eleonora Di Valentino,[2,‡] Rafael C. Nunes,[3,4,§] and Sunny Vagnozzi[5,6,¶]
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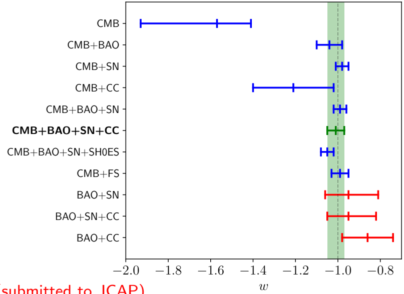

Escamilla *et al.*, arXiv:2307.14802 (submitted to JCAP)

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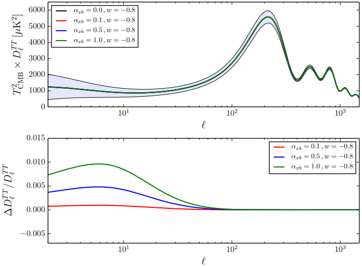
Where to from here? What about the S_8 tension?

Early times: a relatively successful early-time model (EDE and variants, Δm_e,...)
Late times: scattering-type new physics (at 1st order does not affect background but only perturbations) involving DM and/or DE $\rightarrow$ decouple S_8-solving effects from H_0-solving ones, combine the two constructively?

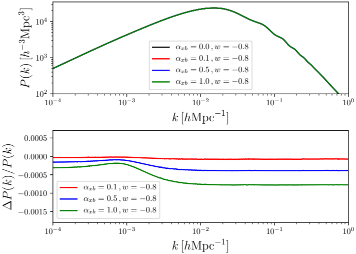
Example: DE-baryon scattering

$$\dot{\theta}_b = -\mathcal{H}\theta_b + c_s^2 k^2 \delta_b + \frac{4\rho_\gamma}{3\rho_b} a n_e \sigma_T (\theta_\gamma - \theta_b) + (1+w_x)\frac{\rho_x}{\rho_b} a n_e \sigma_{xb}(\theta_x - \theta_b)$$

$$\dot{\theta}_x = -\mathcal{H}(1-3c_s^2)\theta_x + \frac{c_s^2 k^2}{1+w_x}\delta_x + a n_e \sigma_{xb}(\theta_b - \theta_x)$$



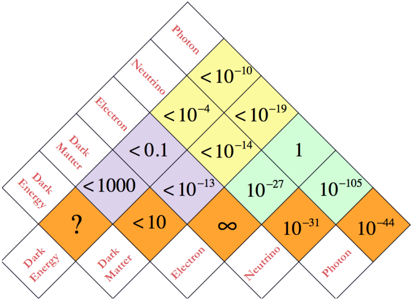
SV *et al.*, MNRAS 493 (2020) 1139



SV *et al.*, MNRAS 493 (2020) 1139

Dark scattering (and S_8)

Lots of room for dark scattering



Simpson, PRD 82 (2010) 083505

Concrete recent example explicitly discussing the S_8 tension

Sigma-8 tension is a drag

Vivian Poulin,[1] José Luis Bernal,[2] Ely D. Kovetz,[3] and Marc Kamionkowski[2]
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(Received 23 September 2022; accepted 8 June 2023; published 30 June 2023)

Poulin *et al.*, PRD 107 (2023) 123538

Possible underlying Lagrangian: "Type 3" coupled DE models (scalar field derivative coupling to velocity)

Models of dark matter coupled to dark energy

A. Pourtsidou, C. Skordis, and E. J. Copeland
Phys. Rev. D **88**, 083505 – Published 9 October 2013

See classification presented in Pourtsidou, Skordis & Copeland, PRD 88 (2013) 083505